PROPERTIES OF POINT MASS LENSES ON A REGULAR POLYGON AND THE PROBLEM OF MAXIMUM NUMBER OF IMAGES

S. Mao

Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Strasse 1, 85740 Garching, Germany

A. O. Petters

Department of Mathematics, Princeton University, Princeton, NJ, 08544-1000, USA

H. J. Witt

Astrophysikalisches Institut Potsdam, An der Sternwarte 16, 14482 Potsdam, Germany

We study the critical curves, caustics, and multiple imaging due to one of the simplest many-body gravitational lens configurations: equal-mass point masses on the vertices of a regular polygon. Some examples of the critical curves and caustics are also displayed. We pose the problem of determining the maximum number of lensed images due to regular-polygon and general point mass configurations. Our numerical simulations suggest a maximum that is linear, rather than quadratic, in the number of point masses.

1 Critical Curves and Caustics

Suppose that g point masses m_k , where $m_k = 1/g$, are on the vertices z_k of a radius r regular polygon centered at the origin: $z_k = r e^{i2\pi(k-1)/g}$, where k = 1, ..., k and $g \ge 2$. [A point mass lens (g = 1) produces a circular critical curve and point caustic; light sources off the point caustic have two lensed images.] The associated lens equation, expressed in complex quantities (Bourassa, Kantowski & Norton, ¹ Witt ²), is given by

$$z_s = z - \sum_{k=1}^g \frac{m_k}{\bar{z} - \bar{z_k}} = z - \frac{\bar{z}^{g-1}}{\bar{z}^g - r^g},$$
 (1)

where z_s is the light-source position. The lens equation defines a mapping, $\eta: z \mapsto z_s$, from the lens plane into the light source plane. Lensed images of a light source at z_s are solutions z in \mathbf{C} of the lens equation. Critical curves (i.e., set of all infinitely magnified lensed images) are given by setting the Jacobian determinant J of η equal to zero:

$$J = 1 - \frac{\partial z_s}{\partial \bar{z}} \frac{\overline{\partial z_s}}{\partial \bar{z}} = 0.$$

This is solved by $\frac{\partial z_s}{\partial \bar{z}} = e^{i\phi}$, where $0 \le \phi < 2\pi$. The critical curves are then the solution curves $z(\phi)$, where $0 \le \phi < 2\pi$, of

$$p(z) = z^{2g} + e^{i\phi}z^{2g-2} - 2r^gz^g + (g-1)r^ge^{i\phi}z^{g-2} + r^{2g} = 0.$$
 (2)

Since z = 0 is not a root of Eq.(2), no critical curve and no caustic passes through the origin. Caustics (i.e., set of positions from which a light source has at least one

infinitely magnified lens image) are the η -images of critical curves. By Eq.(2), there are at most 2g critical curves; hence, the same for caustics. If critical curves merge for a given parameter value, then p(z) has a double or higher order zero. But $\frac{\partial^2 z_s}{\partial z^2}$ is equivalent to a complex polynomial in z of degrees 3 and 6 for g=2 and 3,resp., and degree 2g+1 for $g\geq 4$. In the latter case, one solution is z=0, which cannot lie on no critical curve or caustic. Thus, if g=2,3, then there are at most 3,6 beak-to-beak caustics, resp., while at most 2g occur for $g\geq 4$. A general g-point mass system has at most 3g-3 beak-to-beaks (Witt & Petters³).

Examples of the critical curves and caustics are shown in Figure 1 as a function of r for the case g=6.

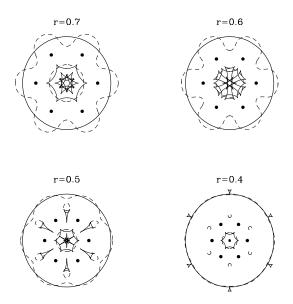


Figure 1: Critical curves (dashed lines) and caustics (thin solid lines) are shown for equal-mass sextuple lenses (filled dots) as a function of the radius of the regular polygon. A unit circle is indicated with the thick solid line.

2 Number of Lensed Images of a Light Source at the Center

For a light source at the center of the polygon $(z_s = 0)$, the lens equation becomes

$$\rho = \frac{\rho^{g-1}}{\rho^g - r^g e^{ig\theta}},\tag{3}$$

where we expressed z in the polar coordinates $\rho e^{i\theta}$. An immediate solution is $\rho = 0$. For any other solution, $e^{ig\theta}$ must be real and, hence, is either +1 or

-1. Eq.(3) can be recast into $f_{\pm}(\rho) \equiv \rho^g - \rho^{g-2} \pm r^g = 0$, where f_+ and f_- correspond to the lens equation with $e^{i\theta} = -1$ and +1, resp. If n_{\pm} are the number of positive zeros of f_{\pm} , then the total number of images is then simply given by, $N = g(n_+ + n_-) + 1$, where the factor of g arises due to rotational symmetry. If g = 2, then $n_+ = 0,1$ for $r \geq 1$ and r < 1, resp., and $n_- = 1$. It follows that N = 3 or 5 if g = 2. Now, suppose that $g \geq 3$. By Descartes' rule of signs, we have $n_{\pm} \leq \#(\text{sign changes in coefficients of } f_{\pm})$. Consequently, $n_+ \leq 2$ and $n_- \leq 1$ for $g \geq 3$. Hence, $N \leq 3g + 1$. This upper bound is also the maximum, that is, it is achievable for each g. In fact, let $r_{cr} = \left(\rho_m^{g-2} - \rho_m^g\right)^{1/g}$, where $\rho_m = [(g-2)/g]^{1/2}$. It can be shown that if $g \geq 3$, then $n_- = 1$ and $n_+ = 0, 1, 2$ for $r > r_{cr}$, $r = r_{cr}$, $r < r_{cr}$, resp. Thus, if $g \geq 3$, then N = g + 1, 2g + 1, 3g + 1 for $r > r_{cr}$, $r = r_{cr}$, and $r < r_{cr}$, resp.

3 Open Problem

The bounds on the total number of lensed images due to g point masses (not necessarily on a regular polygon) are known to be $g+1 \le N \le g^2+1$. The lower bound follows rigorously from Morse theory (Petters ⁴), while the upper bounds can be shown using a trick substitution (Witt — see Ref. 2), or, via resultants (Petters ⁵). The lower bound g+1 is achievable for each g (Petters ⁶); hence, it is the minimum number of lensed images. We do not know whether the upper bound g^2+1 is attainable for each g. For g point masses on the vertices of a regular polygon, and light source not necessarily at the origin, the maximum number of images appears to be 3g+1. Our numerical simulations for generic point-mass configurations seem to confirm this limit as well. It is unknown to the authors whether the maximum number of lensed images due to a general g point mass lens system is linear, i.e., does $N_{max} = gn_1 + n_2$ for each $g \ge 1$, where n_1 and n_2 are fixed positive integers?

Acknowledgments

S.M. is partly supported by the "Sonderforschungsbereich 375-95 für Astro–Teilchenphysik" der Deutschen Forschungsgemeinschaft. A.P. was supported in part by NSF Grant No. DMS-9404522.

References

- R. R. Bourassa, R. Kantowski, and T. D. Norton, Astrophys. J. 185, 747 (1973).
- 2. H. J. Witt, Astron. Astrophys. 236, 311 (1990).
- 3. H. J. Witt and A. O. Petters, J. Math. Phys. 34, 4093 (1993).
- 4. A. O. Petters, J. Math. Phys. 33, 1915 (1992).
- 5. A. O. Petters, J. Math. Phys. 38, 1605 (1997).
- 6. A. O. Petters, Proc. R. Soc. Lond. A 452, 1475 (1996).